

PRESOLAR GRAPHITE FROM THE MURCHISON METEORITE: NOBLE GASES REVISITED. S. Amari¹, R. Gallino², M. Pignatari², ¹Laboratory for Space Sciences and the Physics Department, Washington University, St. Louis, MO 63130, USA (sa@wuphys.wustl.edu), ²Dipartimento di Fisica Generale, Università di Torino, 10125 Torino, Italy (pignatari@ph.unito.it, gallino@ph.unito.it).

Introduction: Ne-E(L) is one of the anomalous noble gas components that ultimately led to the discovery of presolar grains in meteorites. The component is extremely enriched in ²²Ne and this fact prompted the idea that the ²²Ne is from the decay of ²²Na [1]. The carrier of Ne-E(L) is presolar graphite [2], which has been extracted from Murchison [3] and Orgueil [4]. There are four graphite-rich Murchison separates with a range of density (1.6–2.2 g/cm³): KE1 (1.6–2.05 g/cm³), KFA1 (2.05–2.10 g/cm³), KFB1 (2.10–2.15 g/cm³) and KFC1 (2.15–2.20 g/cm³) [3]. KE1 was further separated to improve purity, yielding KE3 (1.65–1.72 g/cm³) [5]. Thus, isotopic features of KE1 and KE3 can be regarded as the same.

One of the most interesting characteristics of presolar graphite is that isotopic features depend on density. Low-density graphite grains, which are characterized by ¹⁸O excesses, Si isotopic anomalies (mainly in the form of ²⁸Si excesses) and high inferred ²⁶Al/²⁷Al ratios (up to 0.1), are believed to have formed in Type II supernovae [6]. On the other hand, high-density graphite grains from KFC1, with their high ¹²C/¹³C ratios and ³⁰Si excesses, most likely formed in low-metallicity AGB (asymptotic giant branch) stars [7]. In this paper, we will reexamine noble gas data previously obtained on the Murchison separates.

Data Sources: Two noble gas analyses were performed: Amari et al. [8] analyzed Ne, Ar, Kr and Xe in bulk (=aggregates) samples from KE1, KFA1, KFB1 and KFC1 by stepwise heating. Nichols et al. [9] analyzed ⁴He and ^{20,22}Ne in single grains from KE3, KFB1 and KFC1.

Discussion: We will mainly focus on the data on KE1 and KE3.

Neon-22. Amari et al. [8] concluded that the ²²Ne in presolar graphite is dominantly from ²²Na, but portion of it originated from the He-shell and that the proportion of ²²Ne of the two origins depend on density. Nichols et al. [9] found 9 ²²Ne-rich grains out of 21 KE3 grains and were able to obtain an upper limit of the ²⁰Ne/²²Ne ratio of 0.01 for grain KE3a-573. Of the 9 grains, 8 grains were previously analyzed for their C, O, and Si isotopic ratios with ion probe. Seven grains show ¹⁸O/¹⁶O ratios higher than the solar ratio (2×10^{-3}), ranging from 6.83×10^{-3} to 0.119. Inferred ⁴⁴Ti/⁴⁸Ti ratios of two of the seven

grains are $(1.06 \pm 0.22) \times 10^{-3}$ and $(2.4 \pm 1.1) \times 10^{-3}$. The grain with the normal ¹⁸O/¹⁶O ratio shows ²⁸Si excess ($\delta^{29}\text{Si} = -235 \pm 97$ ‰, $\delta^{30}\text{Si} = -327 \pm 128$ ‰). All ²²Ne-rich grains show ¹⁸O excesses and/or ²⁸Si excesses, indicating they formed in supernovae.

The lowest ²⁰Ne/²²Ne ratio is found in the He/C zone in supernovae, where ¹⁴N is converted to ²²Ne. The ²⁰Ne/²²Ne ratio in the zone is expected to be 0.096 in a 25M_⊙ model with the solar metallicity by Heger et al. [10] and 0.088 by Chieffi and Limongi [11], which are much higher than an upper limit of the ²⁰Ne/²²Ne ratio of grain KE3a-573. This indicates that the ²²Ne in the grain was originally incorporated as ²²Na, not as ²²Ne implanted onto the grain. If the latter was the case, ²⁰Ne should also have been implanted, resulting in a much higher ²⁰Ne/²²Ne ratio than 0.01.

Although upper limits of ²⁰Ne/²²Ne ratios of the other grains are not available, the portion of ²²Ne from ²²Na can be determined from the Ne isotopic ratios of KE1 (²⁰Ne/²²Ne = 0.0301 ± 0.0018 , ²¹Ne/²²Ne = 0.000118 ± 0.000017). Assuming that the Ne in KE1 is a mixture of ²²Ne from ²²Na, Ne from the He-shell and solar Ne, more than 99% of the ²²Ne in KE1 originated from ²²Na. Thus, all ²²Ne-rich grains contain ²²Ne from ²²Na, as first suggested by Nichols et al. [9]. ²²Na is produced in the O/Ne zone during the hydrostatic burning by ²¹Ne(p,γ)²²Na, where ²¹Ne is produced by ²⁰Ne(n,γ)²¹Ne and protons are produced by ¹²C(¹²C,p)²³Na [11].

Krypton. ⁸⁶Kr/⁸²Kr and ⁸⁰Kr/⁸²Kr are very sensitive to nucleosynthetic conditions. When an unstable isotope ⁸⁵Kr is at the ground state, it decays to ⁸⁵Rb with the half-life of 11 years. When it is at the isomeric state, it decays faster to ⁸⁵Rb ($T_{1/2} = 4.48$ h). During core He burning, the ground and the isomeric states are not thermalized and need to be treated independently. During shell C burning, there is full thermalization between the ground state and the isomeric state [12]. In any case, ⁸⁶Kr yields depend on neutron density. The half-life of ⁷⁹Se strongly depends on temperature. It is much shorter at stellar conditions (one month at $\sim 1 \times 10^9$ K) than in terrestrial conditions (650,000 years) [13]. As a consequence, ⁸⁰Kr yields depend on neutron density and temperature.

Amari et al. [8] have found that the four Murchison separates are enriched s-process Kr (Kr-S)

and that in a $^{86}\text{Kr}/^{82}\text{Kr}$ - $^{83}\text{Kr}/^{82}\text{Kr}$ plot, KE1+KFA1 and KFC1 form two distinct lines, indicating there are two Kr-S components: Kr-SH in KFC1 and Kr-SL in the lower density separate. To infer Kr-SH, we assumed ($^{83}\text{Kr}/^{82}\text{Kr}$)_s = 0.375 instead of 0.30 used by Amari et al. [8], reflecting the improvement of precisions of analyses in neutron capture cross sections (Table 1). Kr-SH, with a high $^{86}\text{Kr}/^{82}\text{Kr}$ ratio, most likely originated from low-metallicity AGB stars as concluded by Amari et al. [8]. Kr-SL was originally associated with high-metallicity AGB stars or a mixture of AGB and massive stars, thus its isotopic ratios were inferred by assuming $^{83}\text{Kr}/^{82}\text{Kr} = 0.30$. Since this value (as well as 0.375) is derived for the main *s*-process component, it is necessary to reevaluate the $^{83}\text{Kr}/^{82}\text{Kr}$ to apply for low-density graphite grains, which formed in supernovae. Since the lowest $^{86}\text{Kr}/^{82}\text{Kr}$ and $^{83}\text{Kr}/^{82}\text{Kr}$ ratios observed in the grains are 1 and 0.7, respectively [8], Kr-SL must have smaller ratios than those.

Table 1. Krypton isotopic ratios

| | 80/82 | 83/82 | 84/82 | 86/82 |
|----------|-----------------|--------|---------------|---------------|
| KE1+KFA1 | 0.070 ±0.045 | ≃0.334 | 1.26 ±0.65 | 0.02 ±0.26 |
| KE1+KFA1 | 0.127 ±0.023 | ≃0.623 | 2.86 ±0.33 | 0.67 ±0.14 |
| KFC1 | 0.030 ±0.047 | ≃0.375 | 2.58 ±0.41 | 4.43 ±0.46 |
| O/C* | 0.405 | 0.334 | 1.41 | 0.0938 |
| O/Ne* | 0.0656 | 0.623 | 1.97 | 0.726 |

* A 25M_⊙ model by [11]

Chieffi and Limongi constructed a set of explosive yields of massive stars of solar metallicity in the mass range of 11–120M_⊙, using the FRANEC code that includes a nuclear network extending to ⁹⁸Mo [11]. Their 15M_⊙ model can be excluded from a source of the Kr in the low-density grains because the $^{86}\text{Kr}/^{82}\text{Kr}$ ratio in the O/Ne zone is 5.05 and the $^{83}\text{Kr}/^{82}\text{Kr}$ ratio of the bottom of the O/C zone is 2.504. In their 25M_⊙ and 35M_⊙ models, the $^{83}\text{Kr}/^{82}\text{Kr}$ ratio is ~0.3 (0.334 in 25M_⊙ and 0.310 for 35M_⊙) in the O/C zone (6.2–6.6M_⊙ in Fig. 1), and ~0.6 (0.623 in 25M_⊙ and 0.597 in 35M_⊙) in the O/Ne zone (2.6–6.2M_⊙ in Fig. 1). The average neutron density in the O/C zone is ≤ 10⁶ n/cm³ [14], whereas that in the O/Ne zone reaches 10¹² n/cm³ [15], which by far exceeds the range of a classical notion of the *s*-process. The KE1+KFA1 data were extrapolated to $^{83}\text{Kr}/^{82}\text{Kr} = 0.334$ and 0.623 to examine the zone where the Kr in the grains originated (Table 1). When extrapolated to 0.334 (hence assuming the Kr was

produced in the O/C zone), both $^{80}\text{Kr}/^{82}\text{Kr}$ and $^{86}\text{Kr}/^{82}\text{Kr}$ ratios are close to zero, whereas the model predicts a much higher $^{80}\text{Kr}/^{82}\text{Kr}$ ratio. When extrapolated to 0.623, the $^{80}\text{Kr}/^{82}\text{Kr}$ ratio from the grains is still higher than the ratio from the model.

It is difficult to conclude which zone is responsible for the Kr in the grains because of huge uncertainties in ^{86}Kr and ^{80}Kr yields. The neutron capture cross section of ^{85}Kr is theoretically estimated with ~80 % uncertainty [16]. The cross section of ^{79}Se is theoretically determined with a huge uncertainty up to 50 %. Moreover, as the half-life of ^{79}Se strongly depends on temperature, a slight difference in temperature can result in a big difference in ^{80}Kr yields.

Summary: Two stellar sources of ²²Na of Ne-E(L) are novae and supernovae. Supernovae are a dominant stellar source of ²²Na in low-density graphite grains. The Kr in the grains shows a signature of neutron capture in supernovae. Its detailed origin is remained to be seen.

References: [1] Clayton D. D. (1975) *Nature*, 257, 36-37. [2] Amari S. et al. (1990) *Nature*, 345, 238-240. [3] Amari S. et al. (1994) *Geochim. Cosmochim. Acta*, 58, 459-470. [4] Jadhav M. et al. (2005) *LPS XXXVI*, Abstract #1976. [5] Amari S. et al. (1995) *Astrophys. J.*, 447, L147-L150. [6] Travaglio C. et al. (1999) *Astrophys. J.*, 510, 325-354. [7] Amari S. et al. (2004) *Meteorit. Planet. Sci.*, 39, A13. [8] Amari S. et al. (1995) *Geochim. Cosmochim. Acta*, 59, 1411-1426. [9] Nichols R. H., Jr. et al. (2006) *Geochim. Cosmochim. Acta*, submitted. [10] Heger A. et al. (2006) in preparation. [11] Chieffi A. and Limongi M. (2006) in preparation. [12] Ward R. A. et al. (1976) *Astrophys. J. Suppl.*, 31, 33-59. [13] Klay N. and Käppeler F. (1988) *Phys. Rev. C*, 38, 295-306. [14] Raiteri C. M. et al. (1993) *Astrophys. J.*, 419, 207-223. [15] Bisterzo S. et al. (2004) *Mem. Soc. Astro. It.*, 75, 741-748. [16] Bao Z. Y. et al. (2000) *At. Data Nucl. Data Tables*, 76, 70.

